RESULTS AND FUTURE PERSPECTIVES OF BOREXINO

ALESSIO CAMINATA

INFN GENOA (ITALY)

ON BEHALF OF THE BOREXINO COLLABORATION

CNNP2020 | February 24th-28th, 2020

OUTLINE

- The Borexino detector
- Why solar neutrinos are important
- A couple of very busy years:
  - the Borexino solar neutrino analysis
  - some news from the Borexino geoneutrino analysis
- A look into the future: the CNO neutrino quest.

THE BOREXINO EXPERIMENT

- **Main goal**: the detection of low energies solar neutrinos, in particular ^{7}Be neutrinos.

- **Detection method**: elastic scattering of neutrinos on electrons.

$$\nu_x + e \rightarrow \nu_x + e \quad x = e, \mu, \tau$$

- **Location**: LNGS (Laboratori Nazionali del Gran Sasso)

- **Detection medium**: large mass of organic liquid scintillator.

The expected rate of ^{7}Be solar neutrinos in 100 ton of Borexino scintillator is about 50 counts/day which corresponds to 10^{-9} Bq/kg.

Just for comparison, natural water is about 10 Bq/kg in ^{238}U, ^{232}Th and ^{40}K.
Huge effort to achieve extreme high radiopurity levels

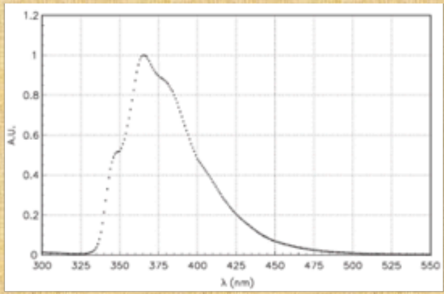
SCINTILLATION LIGHT

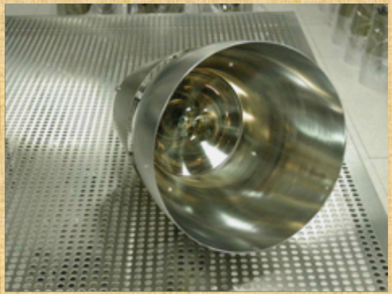
- Reconstruction of:
  - Event position
  - Energy
  - Particle kind (α, β+, β-)

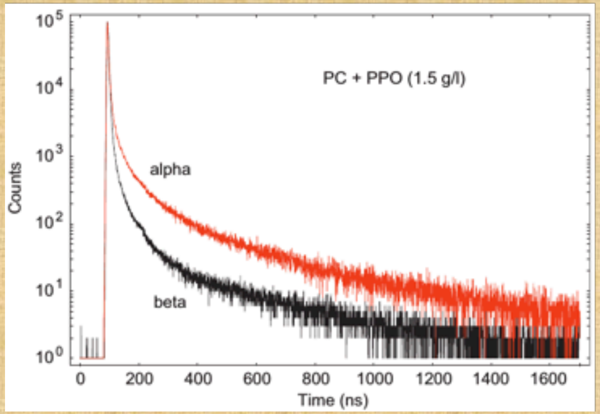
- Isotropic light:
  - No directionality

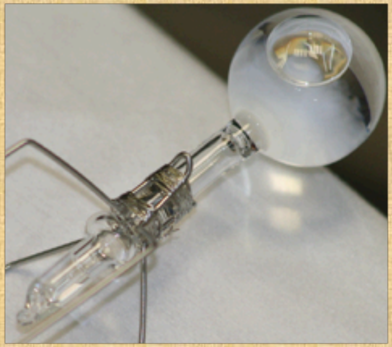
- σ(E) ~ 50 keV
- σ(r) ~10 cm

} @ 1 MeV



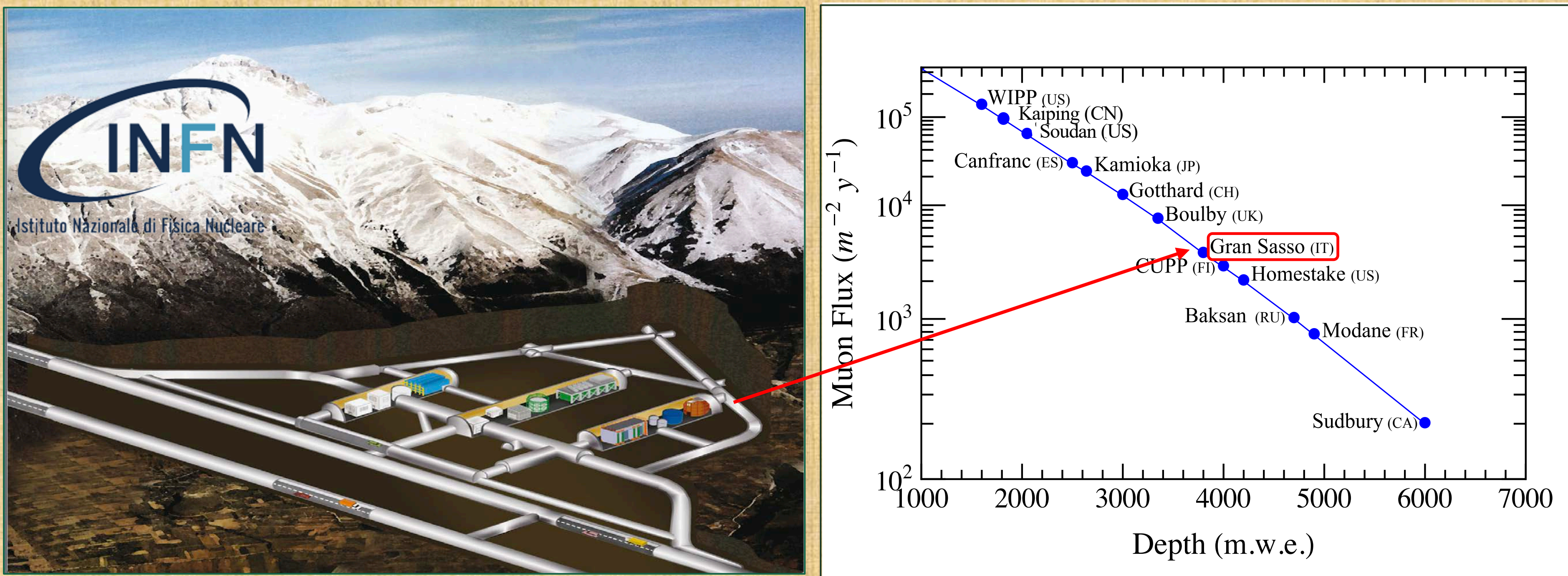






- Performances investigated using calibration sources

Laboratori Nazionali Gran Sasso (ITALY)



The **LNGS** altitude is 963 m and the average rock cover is about 1400 m.

The shielding capacity against cosmic rays is about 3800 m.w.e.:

→ in Borexino the muon flux is reduced by a factor 10^6 with respect to the surface.

$$\Phi(\mu) \sim 1\ \mu/\mathrm{m}^2/\mathrm{h}$$

DETECTOR DESIGN

Scintillator:

280 ton of PC+PPO in a 125 μm thick nylon vessel;

Fiducial mass ~ 100 ton;

Electron density: $(3.307 \pm 0.003) \times 10^{29}/\mathrm{ton}$

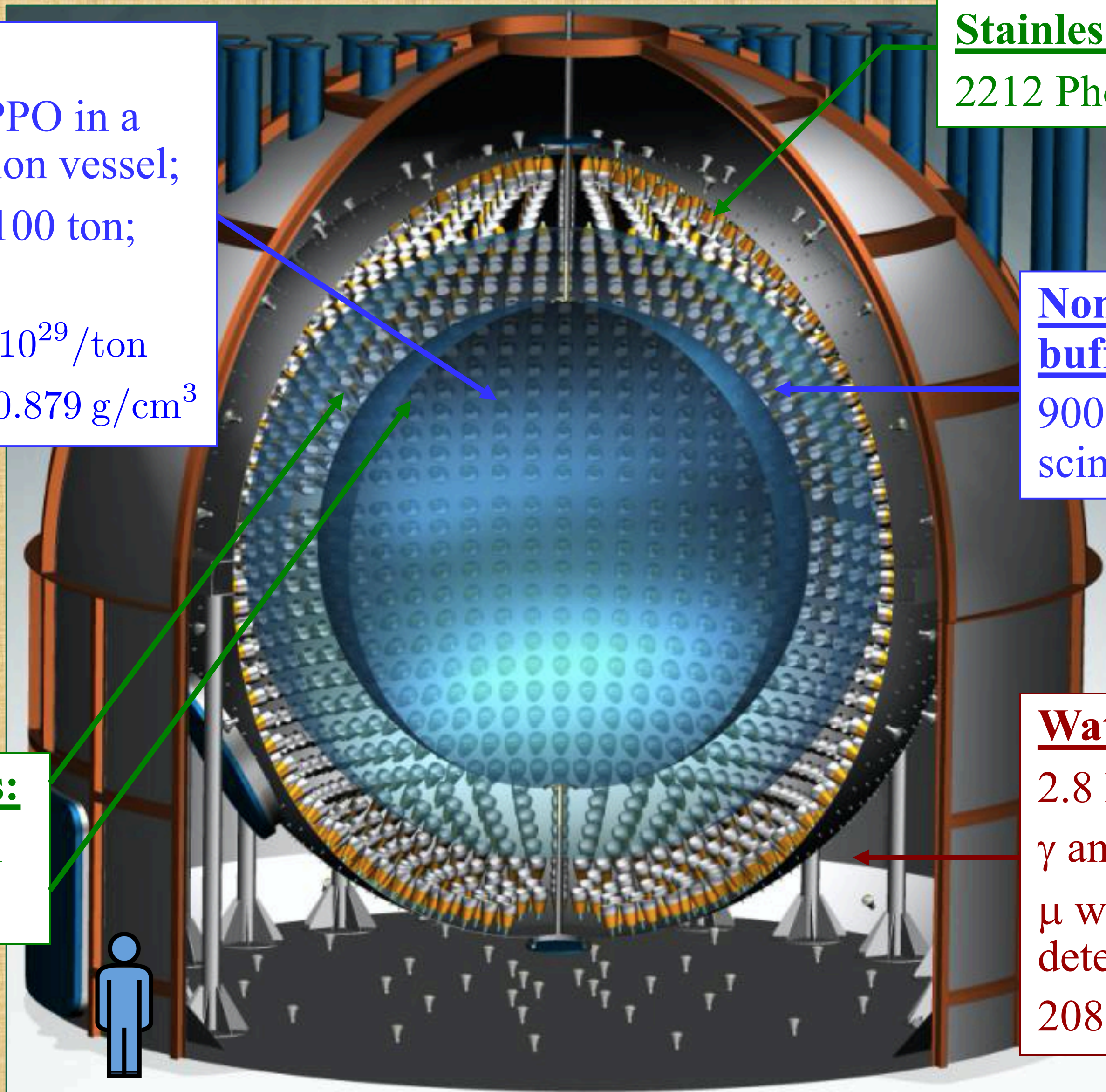
Mass density: $\simeq 0.879\ \mathrm{g/cm^3}$

Stainless Steel Sphere:

2212 Photomultipliers

Non-scintillating buffer:

900 ton of quenched scintillator



Nylon vessels:

Outer: 5.50 m

Inner: 4.25 m

Water Tank:

2.8 kton of pure H_2O

γ and n shield

μ water Čerenkov detector

208 PMTs in water

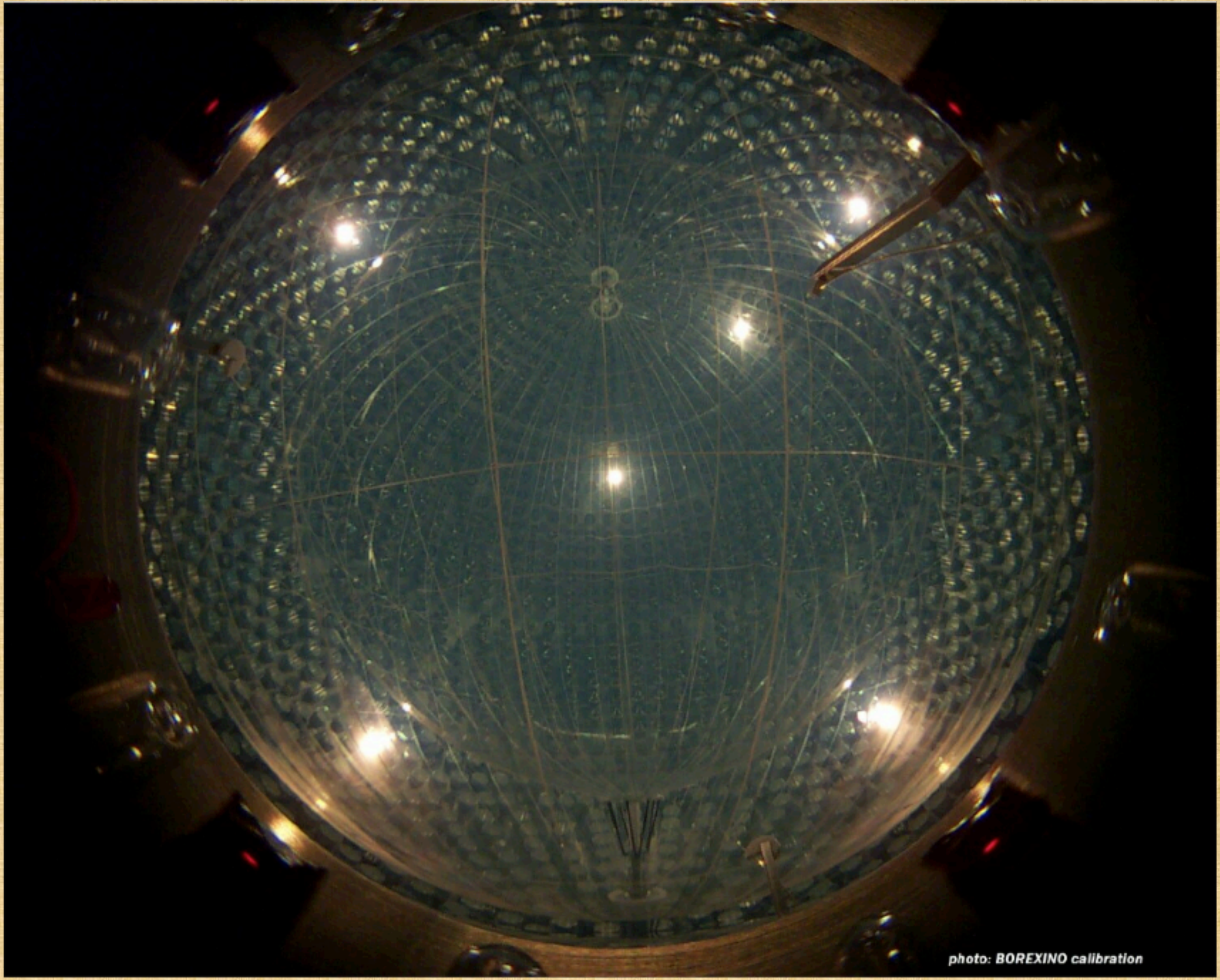
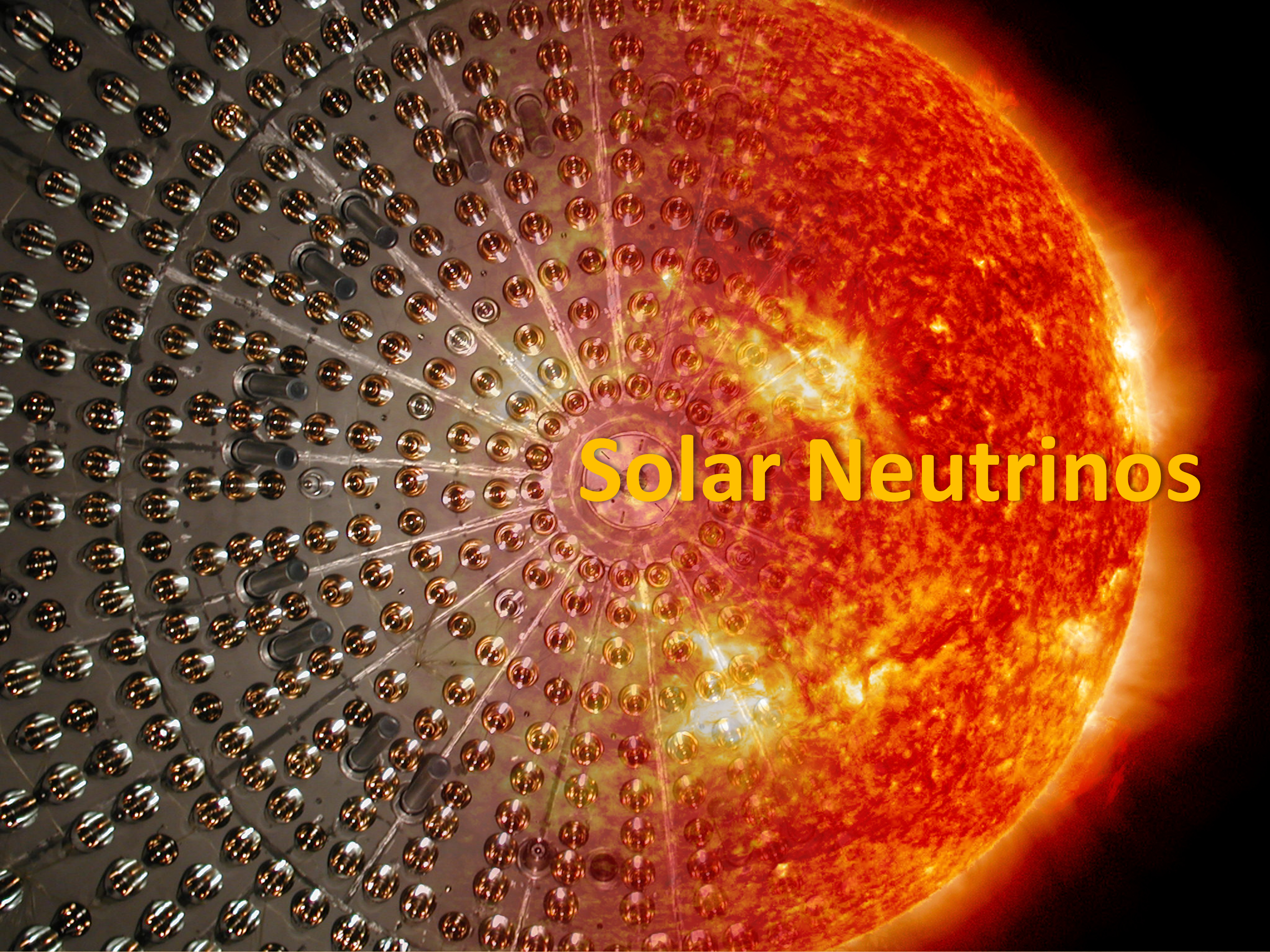


photo: BOREXINO calibration

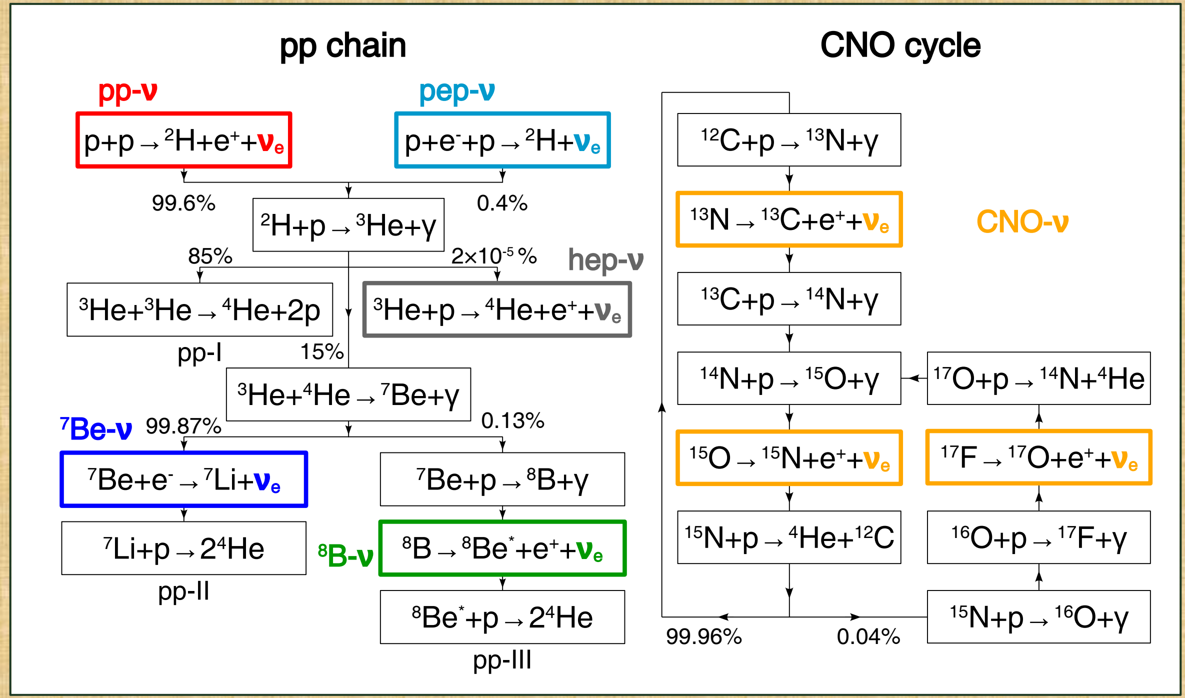

Solar Neutrinos

Studying the Sun with Neutrinos…

$$4\,\mathrm{p} \rightarrow \alpha + 2\,\mathrm{e}^{+} + 2\,\nu_{\mathrm{e}} \qquad \mathrm{E_{released}} \sim 26\,\mathrm{MeV}$$

-While γ massively interact with the solar plasma and take about 10^5 years to reach our star surface, neutrinos stream out the Sun and take just 8 minutes to reach the Earth

-Performing solar neutrino spectroscopy is the only way to get a real-time snap-shot of the nuclear processes inside the Sun



THE STANDARD SOLAR MODEL

A Standard Solar Model (SSM) is a complex container where input parameters (such as Sun Luminosity, Age, Mass, Radius, Chemical elements abundances, Metallicity….) are considered all together and result in expectations about the neutrino fluxes and helioseismology.

The METALLICITY Puzzle

Flux	B16-GS98	B16-AGSS09met
$\Phi(pp)$	$5.98(1 \pm 0.006)$	$6.03(1 \pm 0.005)$
$\Phi(pep)$	$1.44(1 \pm 0.01)$	$1.46(1 \pm 0.009)$
$\Phi(hep)$	$7.98(1 \pm 0.30)$	$8.25(1 \pm 0.30)$
$\Phi(^7Be)$	$4.93(1 \pm 0.06)$	$4.50(1 \pm 0.06)$
$\Phi(^8B)$	$5.46(1 \pm 0.12)$	$4.50(1 \pm 0.12)$
$\Phi(^{13}N)$	$2.78(1 \pm 0.15)$	$2.04(1 \pm 0.14)$
$\Phi(^{15}O)$	$2.05(1 \pm 0.17)$	$1.44(1 \pm 0.16)$
$\Phi(^{17}F)$	$5.29(1 \pm 0.20)$	$3.26(1 \pm 0.18)$

Model and Solar Neutrino Fluxes. Units Are: 10^{10}(pp), 10^9 (^{7}Be), 10^8 (pep, ^{13}N, ^{15}O), 10^6(^{8}B, ^{17}F), and 10^3(hep) $cm^{-2}\,s^{-1}$

→ About 9 % difference (^{7}Be)

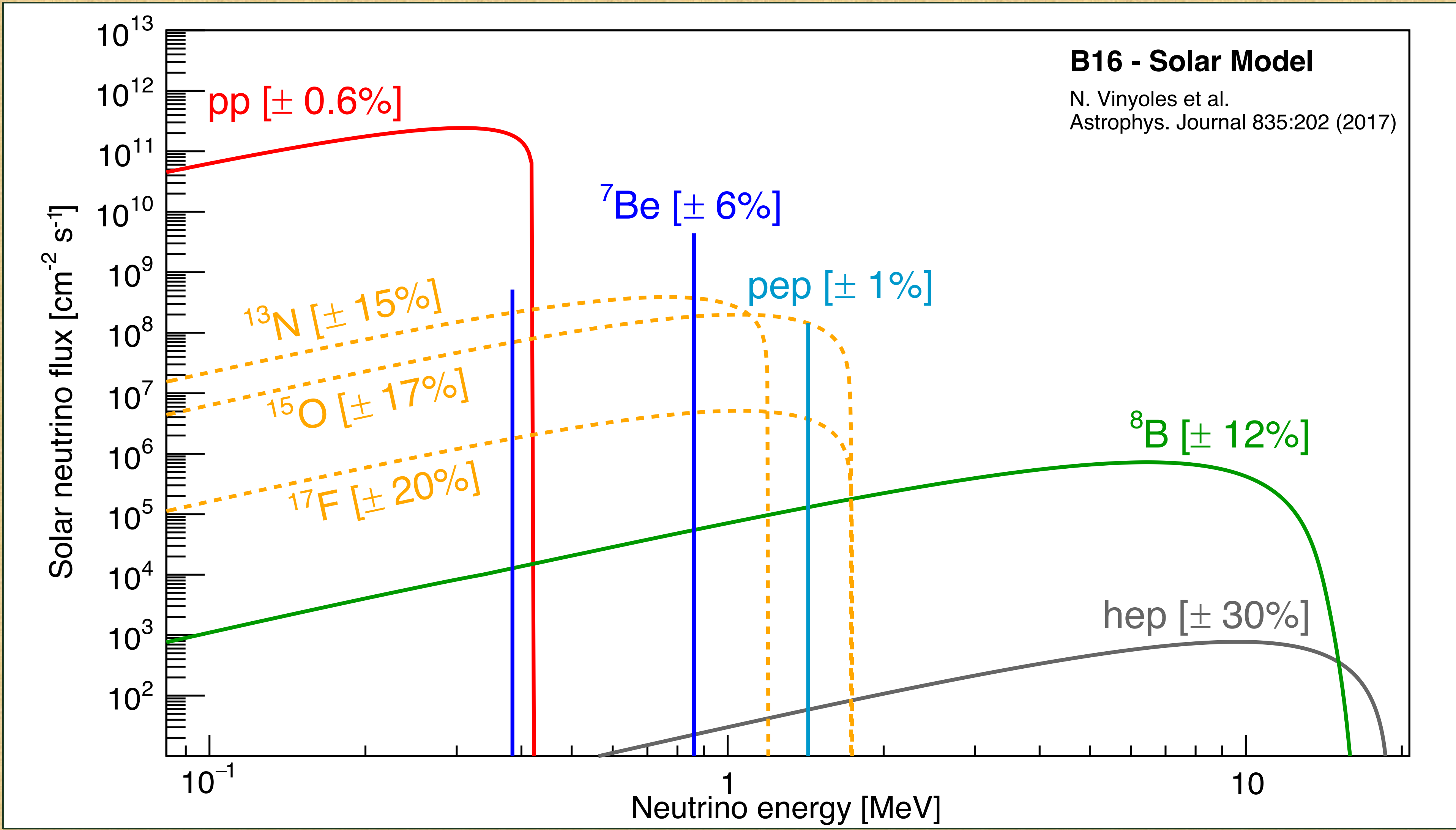
→ About 18 % difference (^{8}B)

} About 28 % difference (^{13}N, ^{15}O, ^{17}F)

B16 - Solar Model

N. Vinyoles et al.
Astrophys. Journal 835:202 (2017)

THE SOLAR NEUTRINO SPECTRUM



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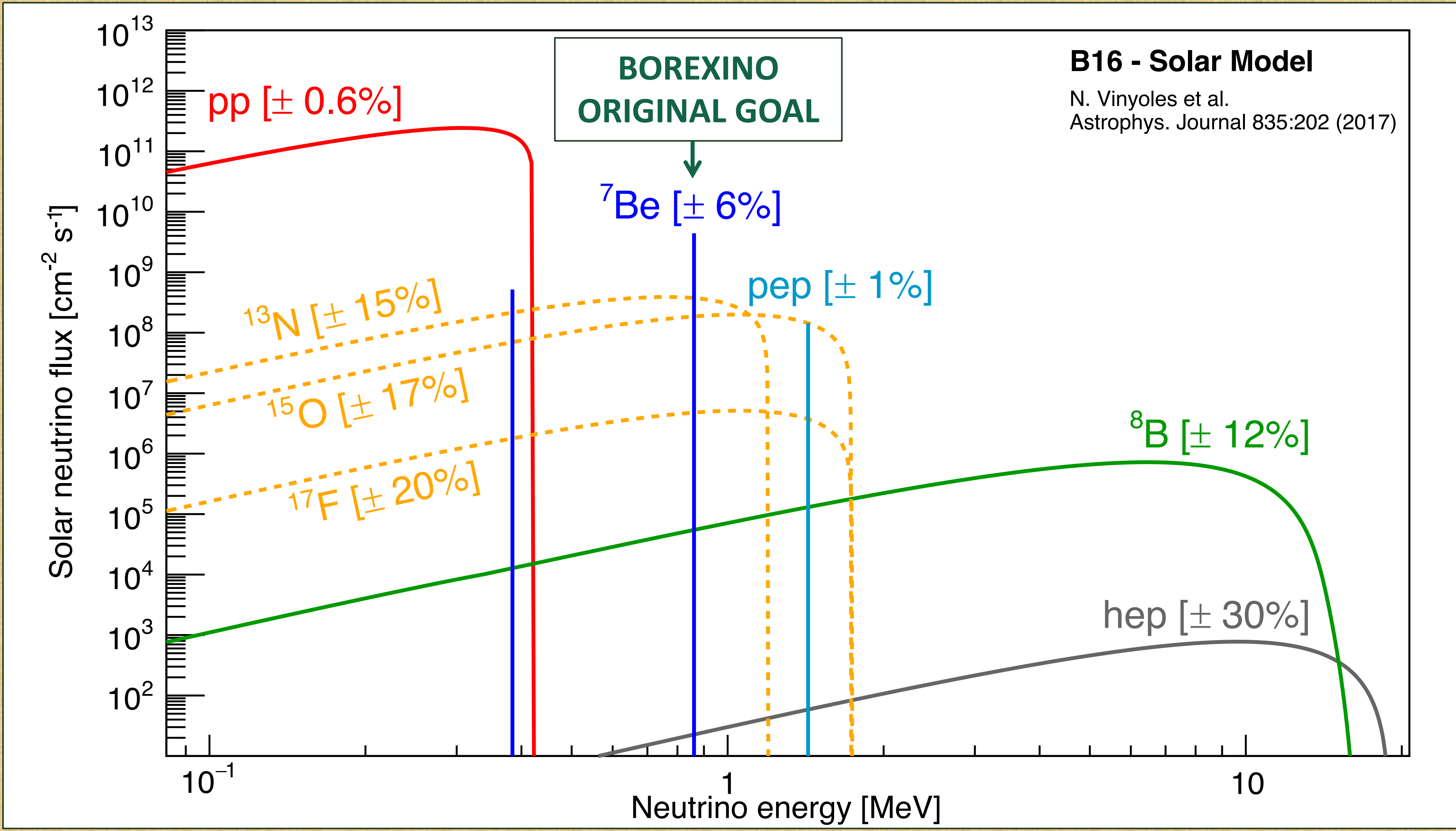

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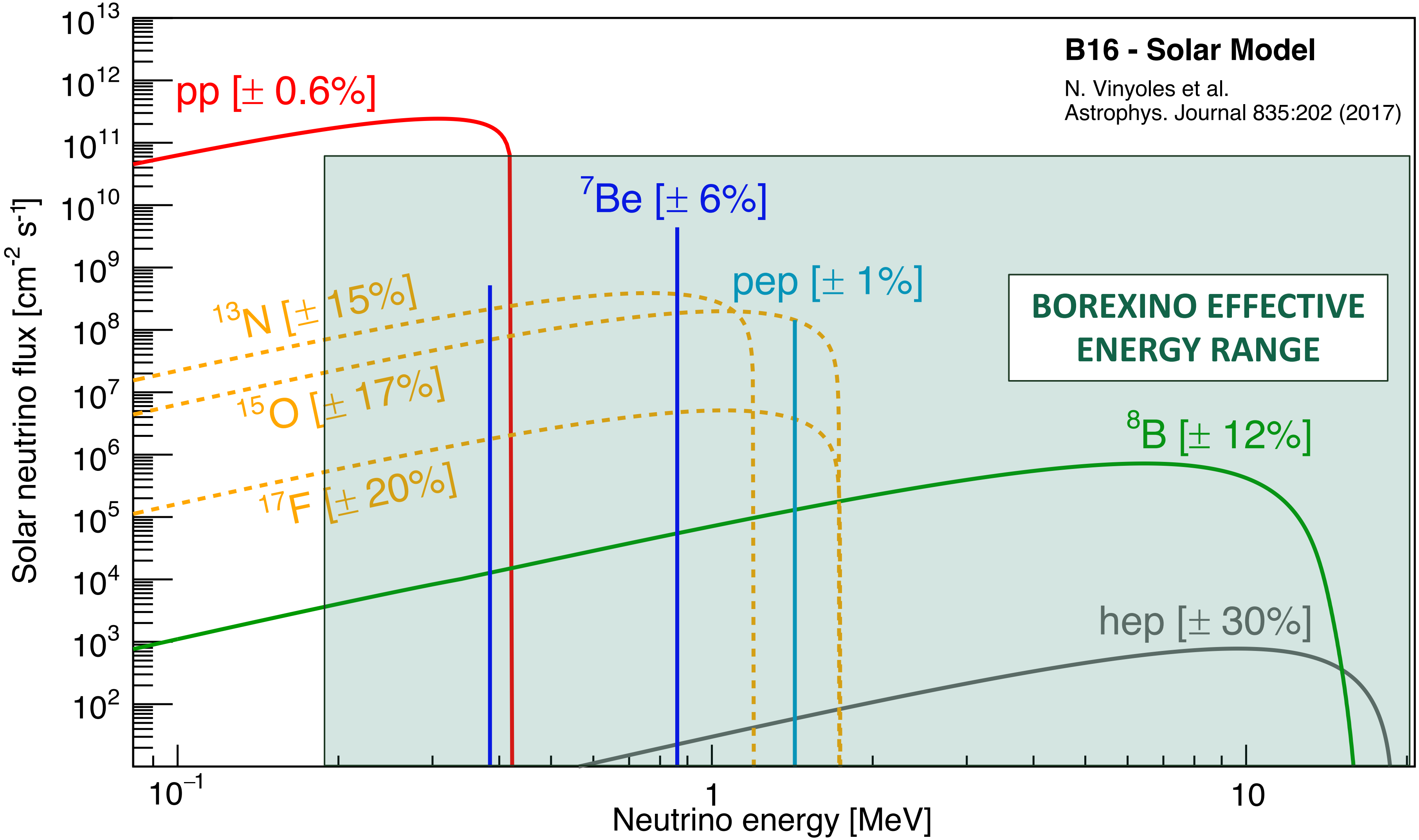


THE BOREXINO RESULTS… SO FAR

Phase I (2007-2010)

2007 — **2010** — **2012** — **2016** — **now**

Direct measurements of

- ^{7}Be flux: 1st observation + precise measurement (5%);
- Absence of day/night asymmetry for ^{7}Be signal
  => MSW-LMA singled out (> 8.5σ);
- ^{8}B flux with low E threshold;
- pep flux: 1st observation;
- CNO upper limit (best to that date).

THE BOREXINO RESULTS... SO FAR

Purification

Phase I (2007-2010)

2007 — 2010 — 2012 — 2016 — now

2011 – 2nd Purification (6 cycles)
Further radiopurity improvement

^{85}Kr: reduced by ∼4.6 factor

^{210}Bi: reduced by ∼2.3 factor

^{238}U: < 9.4×10^{-20} g/g (95% C.L.)

^{232}Th: < 5.7×10^{-19} g/g (95% C.L.)

^{210}Po: reduced by > 10 factor due to natural decay

➡ **The scintillator has never been so clean!**

THE BOREXINO RESULTS… SO FAR

<u>Direct measurements of</u>

- pp flux: 1^{st} direct measurement ;
- Geoneutrinos (> 5σ);
- Electric charge conservation (best limit to date);
- Gamma-ray burst corr.
- ^{7}Be flux seasonal modulation;
- New limit on neutrino magnetic moments;
- Comprehensive measurement of pp-chain solar neutrinos (pep signal > 5σ).

THE BOREXINO RESULTS… SO FAR

Purification — Thermal Insulation

Phase I (2007-2010) | **Phase II (2012-2016)** | **Phase III (2016-)**

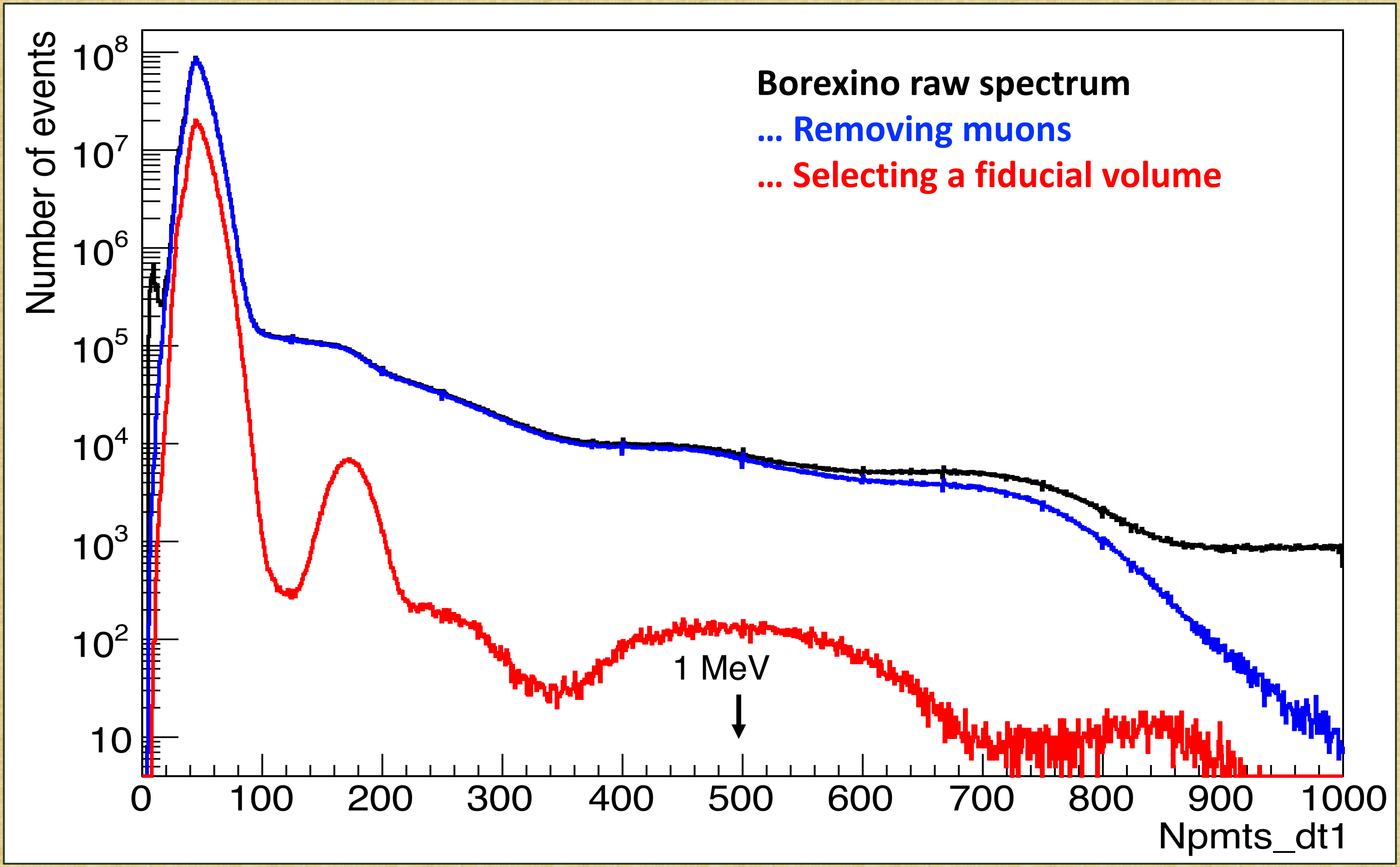
2007 — 2010 — 2012 — 2016 — now

Phase III:

- Geoneutrinos measurement further improved;

➡ **FIRST ATTEMPT TO SEARCH FOR AND MEASURE CNO-ν IS IN PROGRESS NOW!**

HOW TO EXTRACT A NEUTRINO SIGNAL



We need to develop powerful tools to separate the signal from the residual background components

COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

The Borexino experiment has never been so performing…

1. **Improved radiopurity**, because of the purification campaign;
2. **Increased statistics**;
3. **Increased stability** of the detector;
4. **Better comprehension of the details of the energy scale and detector response.**

.... So all challenges at once!

For the first time we are able to perform a simultaneous fit on the whole solar neutrino energy region.

The analysis is carried out on two energy ranges:

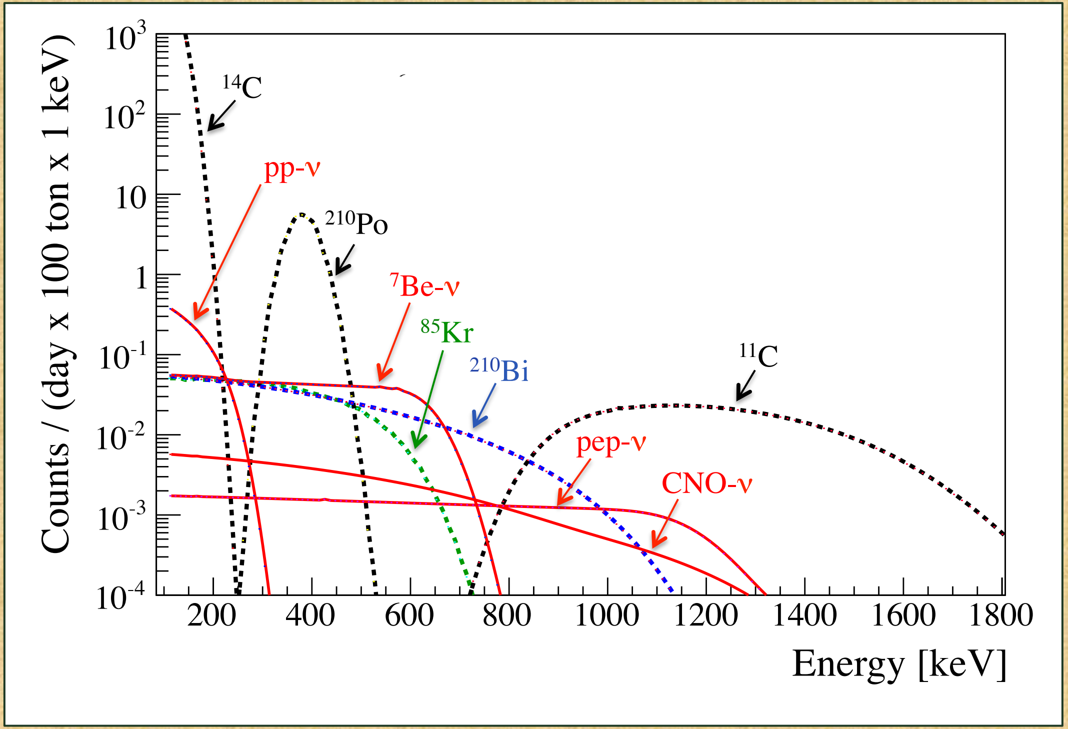
- **LER** (Low Energy Range) between (0.19 –2.93) MeV (pp, pep and ^{7}Be ν)
- **HER** (High Energy Range) between (3.2 -16) MeV (^{8}B and hep ν)

COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

LER analysis *Physical Review D 100, 082004 (2019)*

The analysis follows similar strategies in both the LER and HER but it is differently optimized in the two energy regions to comply with the different backgrounds affecting each specific energy range.



Main LER background sources:

- **^{14}C**: irreducible background in any organic scintillator;
- **^{210}Bi**: comes from ^{210}Pb, is not in equilibrium with the ^{238}U chain;
- **^{210}Po**: comes from ^{210}Bi, is not in equilibrium with the ^{238}U chain;
- **^{85}Kr**: present in air;
- **^{11}C**: produced by μ;
- **pile-up** of events (mainly ^{14}C).

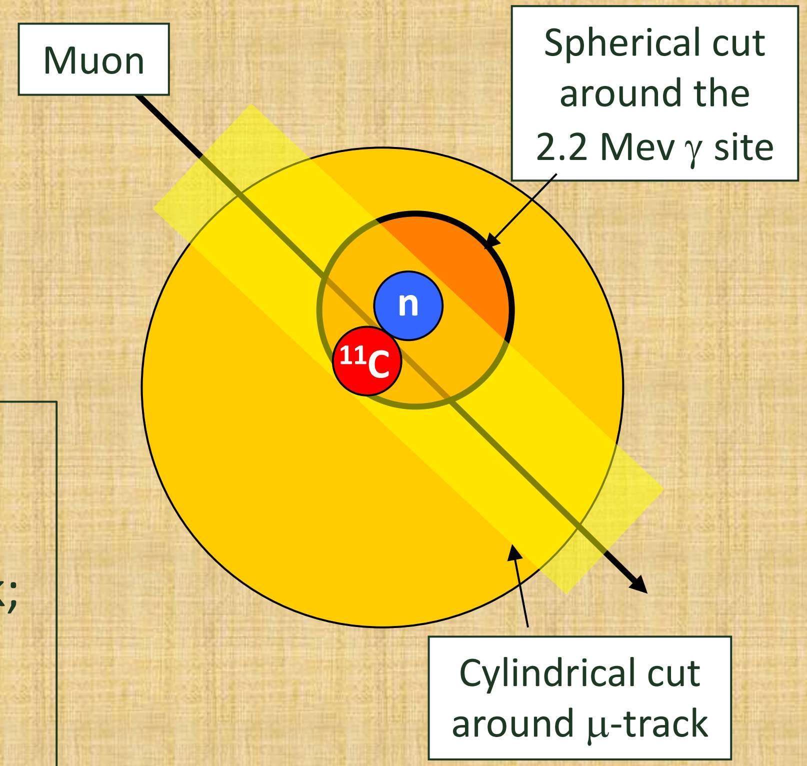
The Three-Fold Coincidence technique (TFC)

The TFC technique is fundamental to improve the fit capability to disentangle the ^{11}C contamination from the pep & CNO neutrino signals.

$$\mu + {}^{12}\mathbf{C} \rightarrow \mu + {}^{11}\mathbf{C} + \mathbf{n}$$

$${}^{11}\mathbf{C} \rightarrow {}^{11}\mathbf{B} + \mathbf{e}^{+} + \nu_{\mathbf{e}}$$
$\tau \sim 30$ min

$$\mathbf{n} + \mathbf{p} \rightarrow \mathbf{d} + \gamma$$
$\tau \sim 260$ μs



The likelihood that a certain event is ^{11}C is obtained using:

- Distance in space and time from the μ-track;
- Distance from the neutron;
- neutron multiplicity;
- Muon dE/dx and number of muon clusters in an event.

COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*
LER analysis *Physical Review D 100, 082004 (2019)*

Main LER analysis features:

- Data-set: December 2011 - May 2016;
- Exposure: 1291.51 days x 71.3 t
  → Exposure 1.6 times the one used for the ^{7}Be measurement @ 5%;
- Fit range: 0.19 – 2.93 MeV.

Two complementary fit methods:

Analytical fit

- model of the detector response;
- possibility to describe unknown time variations;

Monte Carlo fit

- detailed MC modeling tuned on calibrations data;
- sub% accuracy [Astr. Phys. 97 (2018) 136].

The data set is presented as two energy spectra: one with ^{11}C included (TFC-tagged) and one depleted in ^{11}C (TFC-subtracted) which are then simultaneously fit.

The multivariate fit is performed including the likelihood of:

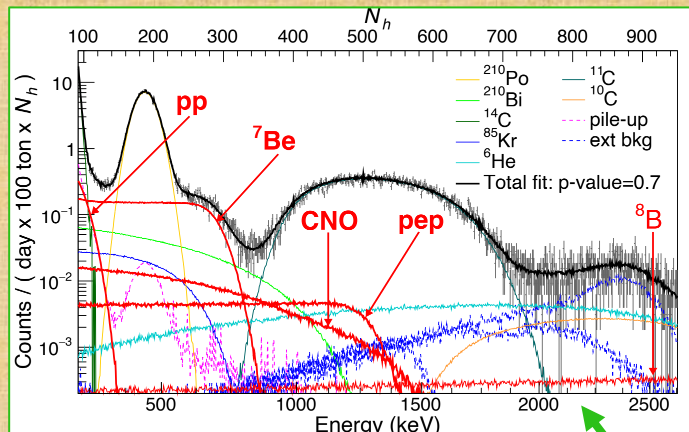
1. Energy spectra (TFC-tagged and TFC-subtracted);
2. e^-/e^+ pulse-shape distribution PS-$\mathcal{L}_{PR}$;
3. Radial distribution.

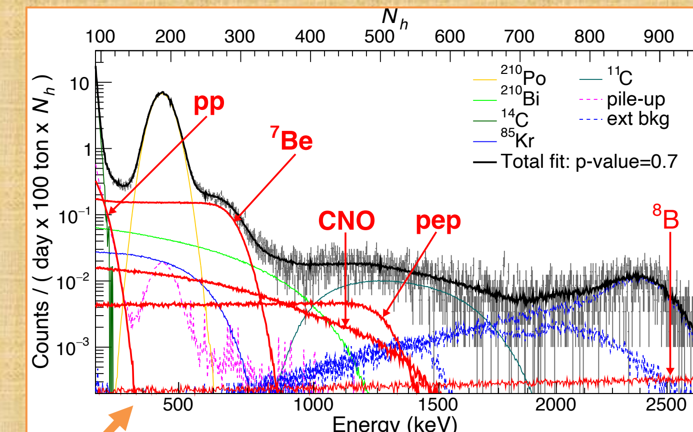
COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

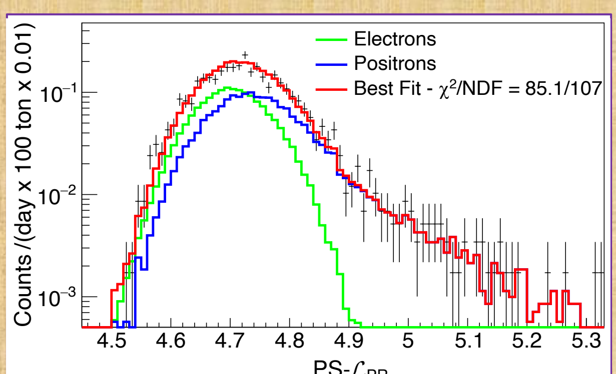
LER analysis *Physical Review D 100, 082004 (2019)*

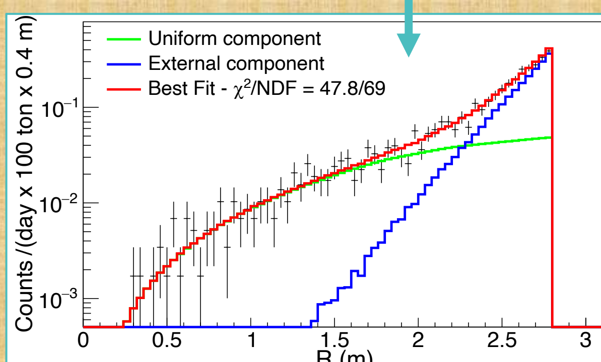
Multivariate fit: neutrino interaction rates obtained by maximizing the binned likelihood function





$$\mathcal{L}_{\rm MV} = \mathcal{L}_{^{11}\rm C-tag} \cdot \mathcal{L}_{^{11}\rm C-sub} \cdot \mathcal{L}_{\rm PS} \cdot \mathcal{L}_{\rm rad}$$





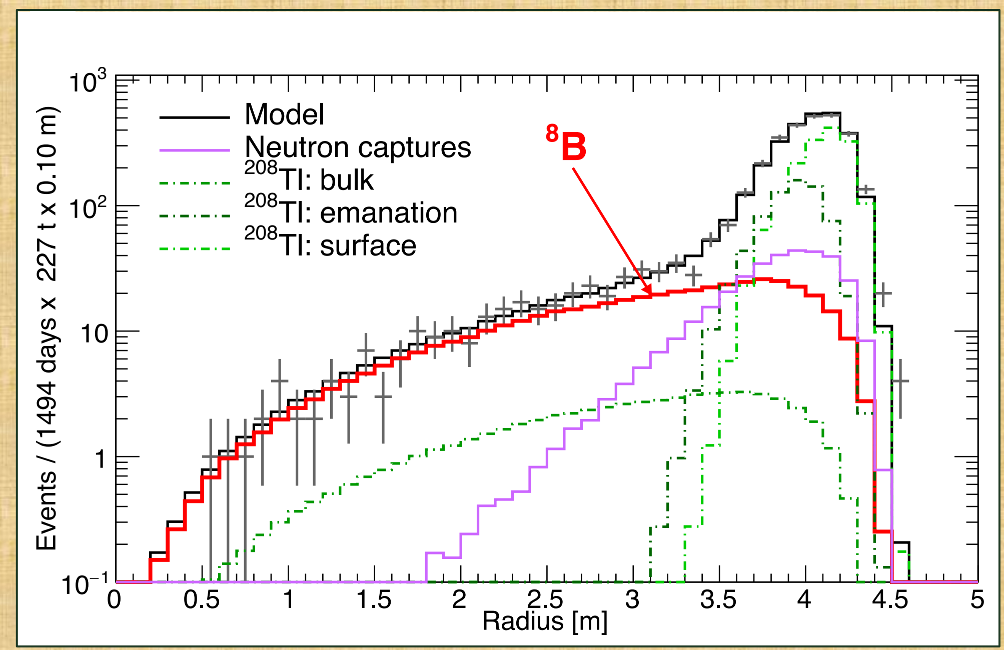
COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*
HER analysis *arXiv: 1709.00756 [hep-ex] (2017)*

Main HER analysis features:

- Data-set: January 2008 - December 2016 (purification period excluded);
- Fiducial mass: extended to the entire active mass (from about 100 t to 300 t);
- Fit range: 3.2 -16 MeV;
- Total exposure: 1.5 kton x year (11.5-fold increase).

New strategy! A MonteCarlo radial fit on Low Energy (HER-I: 3.2-5 MeV) and High Energy (HER-II: 5-16 MeV) sectors so to better handling the background.



Extracting the neutrino signal from data:

Residual backgrounds affecting the ^{8}B energy region are:

- ^{208}Tl (emanated from PMTs, from the vessel or internal);
- cosmogenic isotopes;
- ^{214}Bi (internal).

COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

LER analysis *Physical Review D 100, 082004 (2019)*

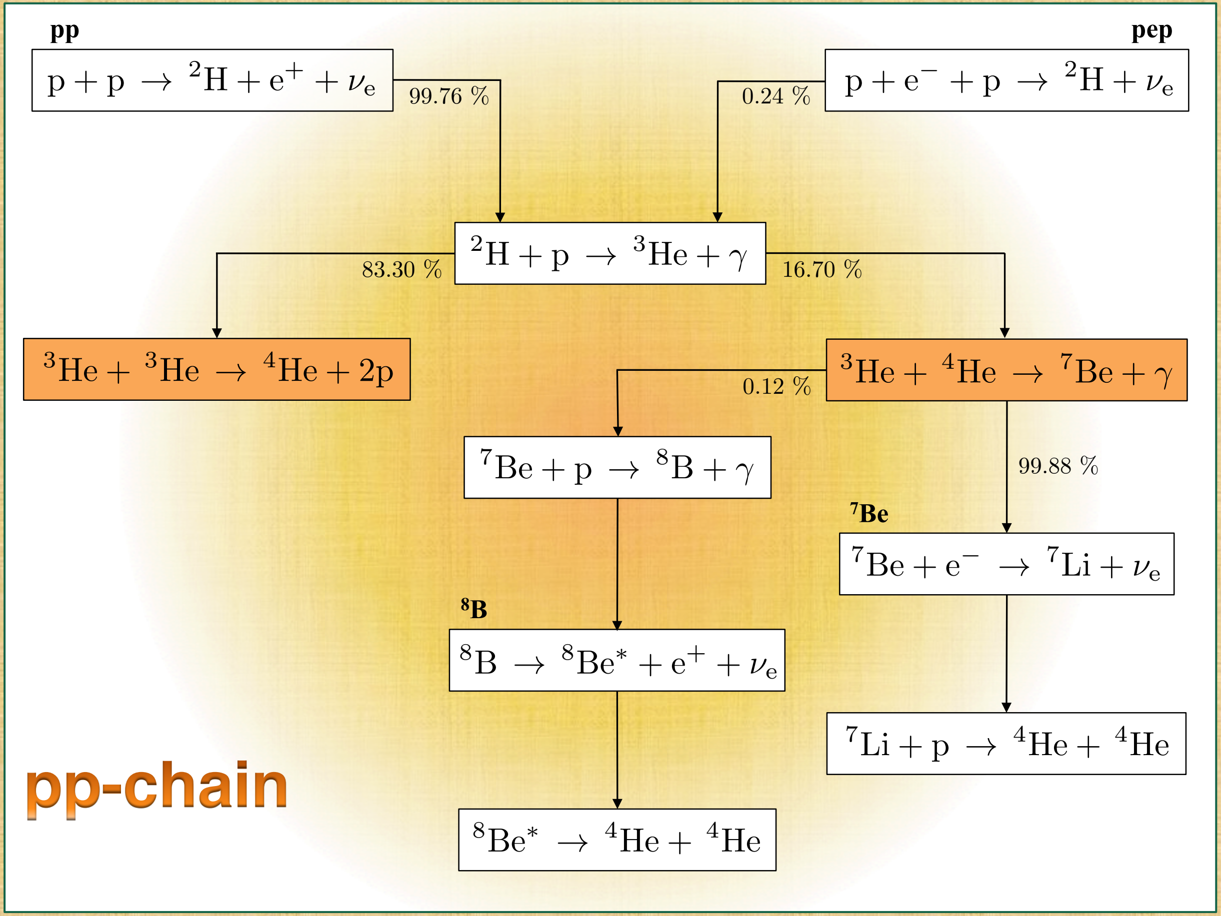
HER analysis *arXiv: 1709.00756 [hep-ex] (2017)*

Solar neutrino	Rate (counts per day per 100 t)	Flux ($cm^{-2}\,s^{-1}$)	Flux–SSM predictions ($cm^{-2}\,s^{-1}$)
pp	$134 \pm 10^{+6}_{-10}$	$(6.1 \pm 0.5^{+0.3}_{-0.5}) \times 10^{10}$	$5.98(1.0 \pm 0.006) \times 10^{10}$ (HZ) $6.03(1.0 \pm 0.005) \times 10^{10}$ (LZ)
^{7}Be	$48.3 \pm 1.1^{+0.4}_{-0.7}$	$(4.99 \pm 0.11^{+0.06}_{-0.08}) \times 10^{9}$	$4.93(1.0 \pm 0.06) \times 10^{9}$ (HZ) $4.50(1.0 \pm 0.06) \times 10^{9}$ (LZ)
pep (HZ)	$2.43 \pm 0.36^{+0.15}_{-0.22}$	$(1.27 \pm 0.19^{+0.08}_{-0.12}) \times 10^{8}$	$1.44(1.0 \pm 0.01) \times 10^{8}$ (HZ) $1.46(1.0 \pm 0.009) \times 10^{8}$ (LZ)
pep (LZ)	$2.65 \pm 0.36^{+0.15}_{-0.24}$	$(1.39 \pm 0.19^{+0.08}_{-0.13}) \times 10^{8}$	$1.44(1.0 \pm 0.01) \times 10^{8}$ (HZ) $1.46(1.0 \pm 0.009) \times 10^{8}$ (LZ)
$^{8}B_{HER}$	$0.223^{+0.015}_{-0.016}{}^{+0.006}_{-0.006}$	$(5.68^{+0.39}_{-0.41}{}^{+0.03}_{-0.03}) \times 10^{6}$	$5.46(1.0 \pm 0.12) \times 10^{6}$ (HZ) $4.50(1.0 \pm 0.12) \times 10^{6}$ (LZ)
CNO	<8.1 (95% C.L.)	$<7.9 \times 10^{8}$ (95% C.L.)	$4.88(1.0 \pm 0.11) \times 10^{8}$ (HZ) $3.51(1.0 \pm 0.10) \times 10^{8}$ (LZ)
hep	<0.002 (90% C.L.)	$<2.2 \times 10^{5}$ (90% C.L.)	$7.98(1.0 \pm 0.30) \times 10^{3}$ (HZ) $8.25(1.0 \pm 0.12) \times 10^{3}$ (LZ)

Comprehensive solar neutrino spectroscopy

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

Astrophysical implications of the results: probing solar fusion



Probing solar fusion by studying the two primary modes of terminating the pp-chain.

$$\mathcal{R} = \frac{2\Phi(^7\mathrm{Be})}{[\Phi(pp) - \Phi(^7\mathrm{Be})]}$$

B16-SSM expected values:

$\mathcal{R} = 0.180 \pm 0.011$ (HZ)

$\mathcal{R} = 0.161 \pm 0.010$ (LZ)

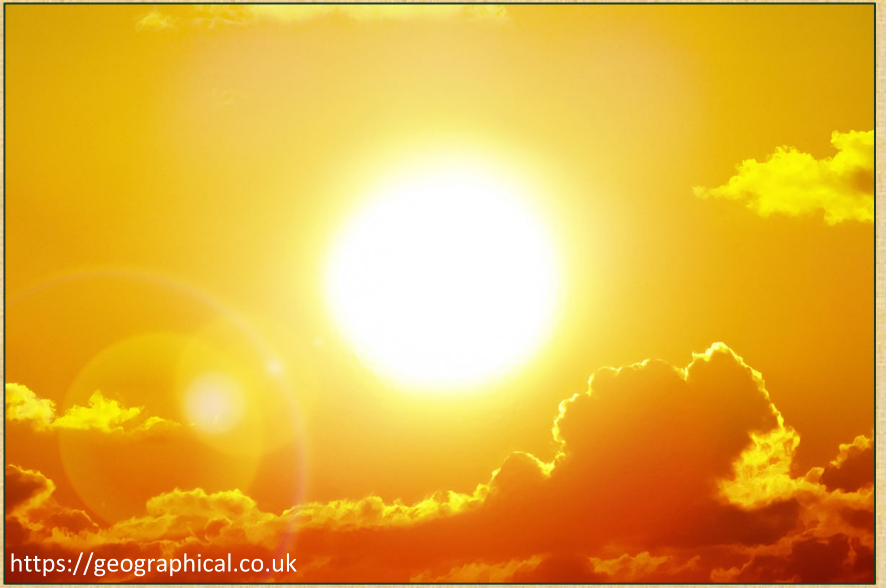
Borexino result:

$\mathcal{R} = 0.178^{+0.027}_{-0.023}$

COMPREHENSIVE SOLAR NEUTRINO SPECTROSCOPY

"Comprehensive measurement of pp-chain solar neutrinos", *Nature 562 (2018) 505*

Astrophysical implications of the results: solar luminosity



Using Borexino results only we can calculate the neutrino solar luminosity:

$$\mathbf{L}_{\nu} = (\mathbf{3.89}\,^{+0.35}_{-0.42}) \times \mathbf{10^{33}\,erg\,s^{-1}}$$

which agrees with the well measured photon value:

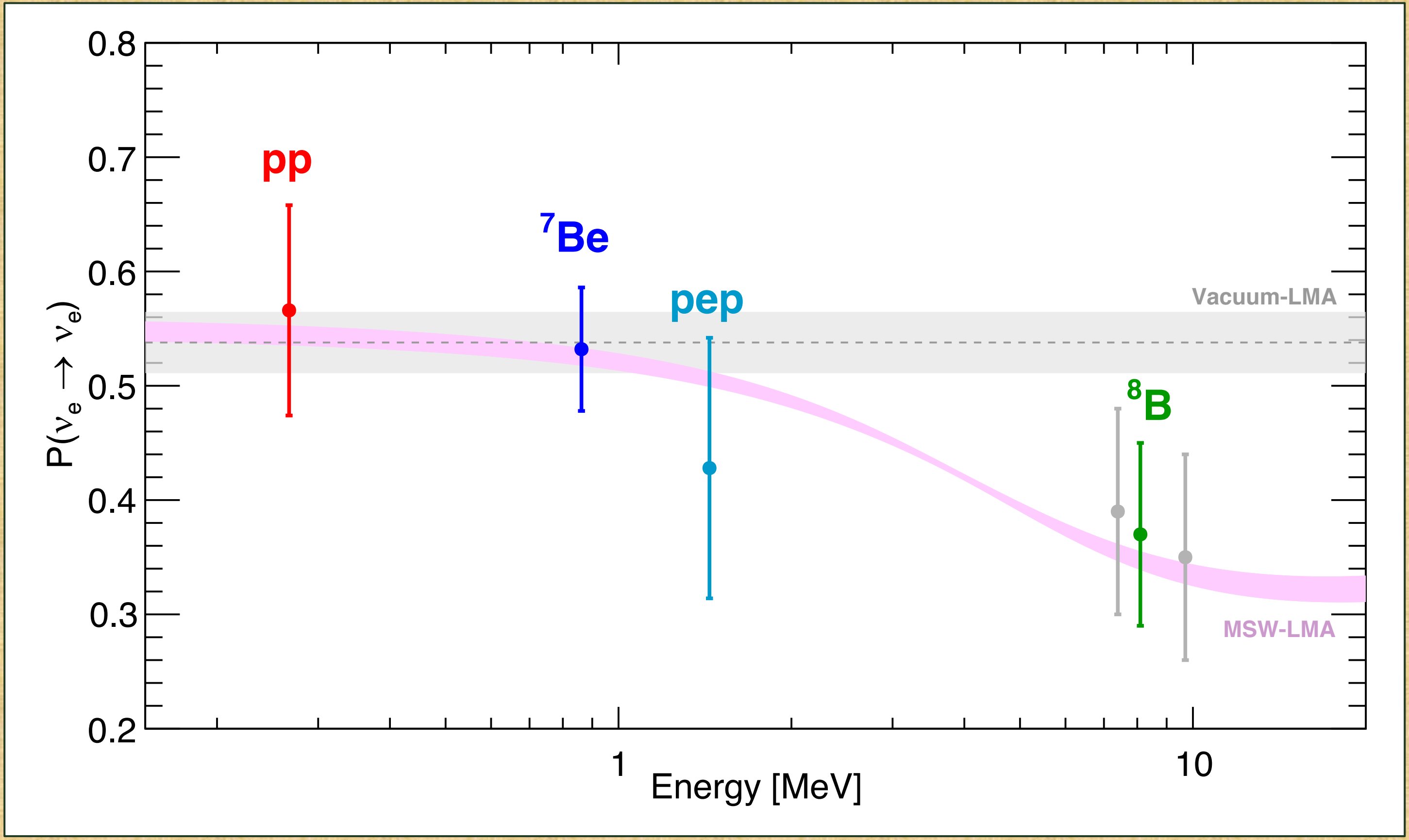
$$\mathbf{L_{ph}} = (\mathbf{3.846 \pm 0.015}) \times \mathbf{10^{33}\,erg\,s^{-1}}$$

→ This confirms the nuclear origin of the solar power!

→ It proves that the Sun has been in thermodynamic equilibrium over the last 10^5 years

STUDYING THE SUN WITH NEUTRINOS… …STUDYING NEUTRINOS WITH THE SUN

Neutrino physics implications of the results: testing MSW-LMA



SSM-HZ solar-ν fluxes from *N. Vinyoles et al., Astrophys. Journal 835:202 (2017)*
Neutrino oscillation parameters from *I. Esteban et al., JHEP 01 (2017).*


Geoneutrinos

WHY TO STUDY GEONEUTRINOS?

Geo-neutrinos are the anti-neutrinos produced in the decays of the progenies of Uranium, of Thorium, and Potassium inside the Earth.

Geo-neutrinos bring to the surface information from the whole planet:
they are a unique direct probe of our Earth's interior!

Open questions:
- What is the radiogenic contribution to the terrestrial heat?
- What is the distribution of the radiogenic elements within the Earth?

GEONEUTRINOS IN BOREXINO

Physical Review D 101, 012009 (2020)

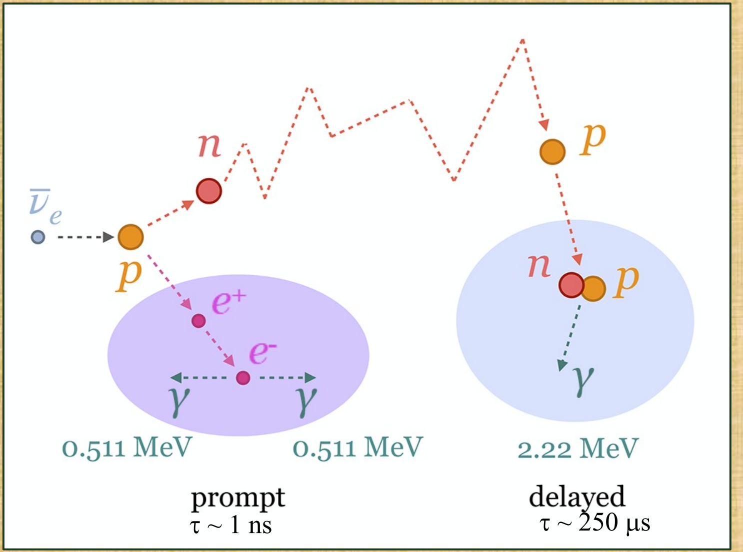
In Borexino, electron anti-ν are detected via the inverse beta decay reaction:

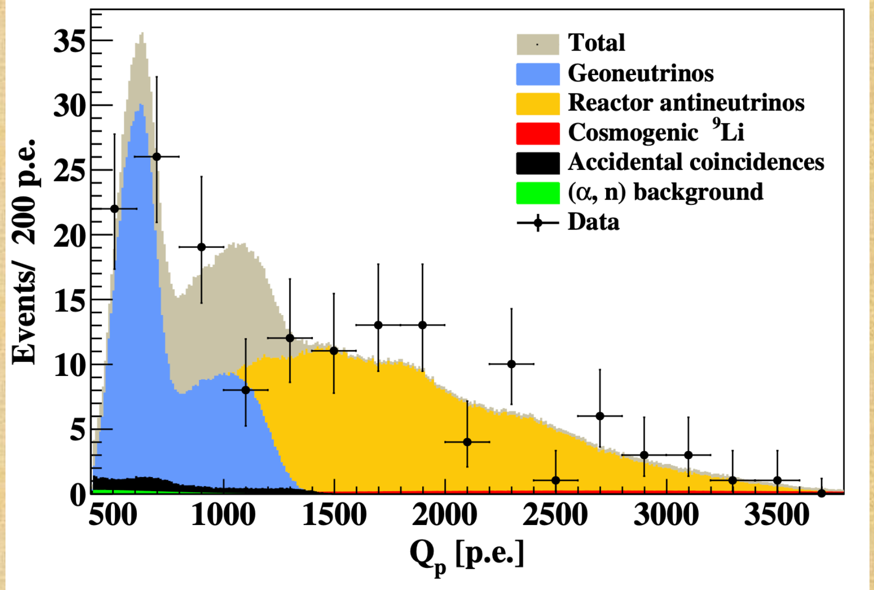
$$\overline{\nu_e} + p \rightarrow e^+ + n$$

$$\sigma_{IBD} \approx 10^{-42}\,\mathrm{cm}^2$$

Charge current interaction → detection of $\overline{\nu_e}$ only

Threshold: 1.8 MeV





- A **total uncertainty** of around ~18% achieved in the geoneutrino signal.
- **Rejection of the null-hypothesis** of mantle signal at 99% C.L.
- **Radiogenic heat**: 2.4σ tension with models predicting lowest amount of mantle signal.

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Physical Review D 101, 012009 (2020)

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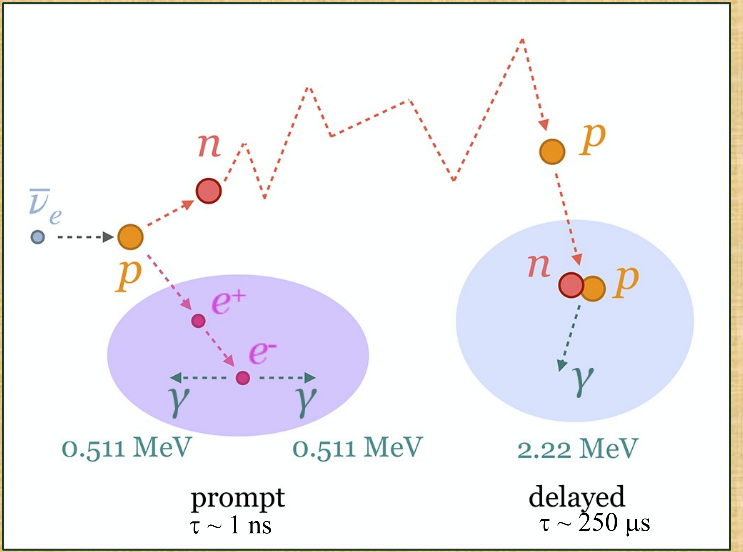

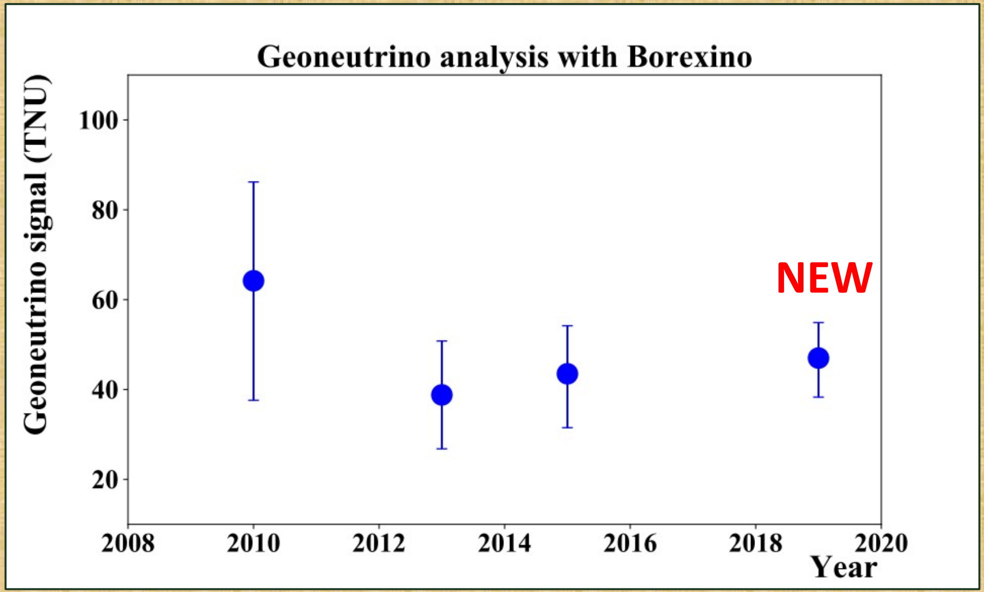

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Other Recent Results from Borexino

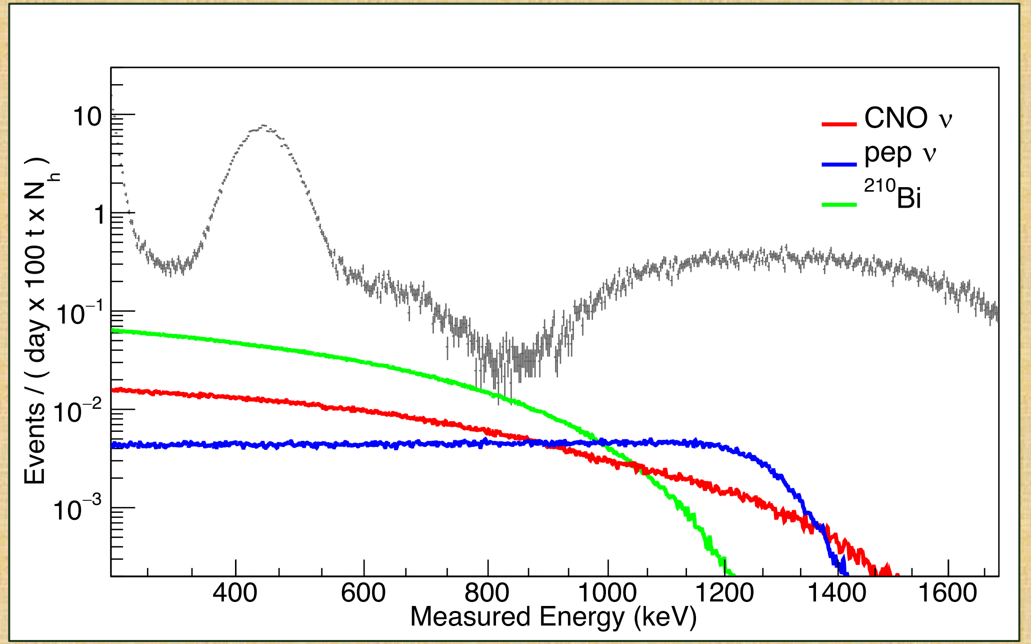
- **Constraints on Flavor-Diagonal Non-Standard Neutrino Interactions from Borexino Phase-II,** ***JHEP 2020, 38***
- **Search for low-energy neutrinos from astrophysical sources with Borexino,** ***arXiv:1909.02422***
- Modulations of the cosmic muon signal in ten years of Borexino data, *JCAP 2019, Feb. 2019*
- Seasonal modulation of the 7Be solar neutrino rate in Borexino , *Astr. Phys. 92 (2017) 21*

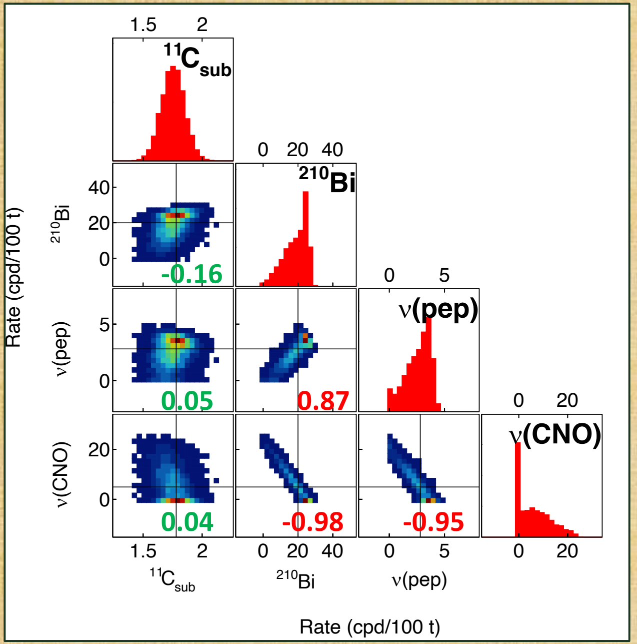
Presented last
Tuesday
by A. Vishneva


The Future
NEXT EXIT

TOWARDS THE CNO-ν MEASUREMENT

The similarity between the CNO, pep and ^{210}Bi spectral shapes limits the sensitivity of Borexino.



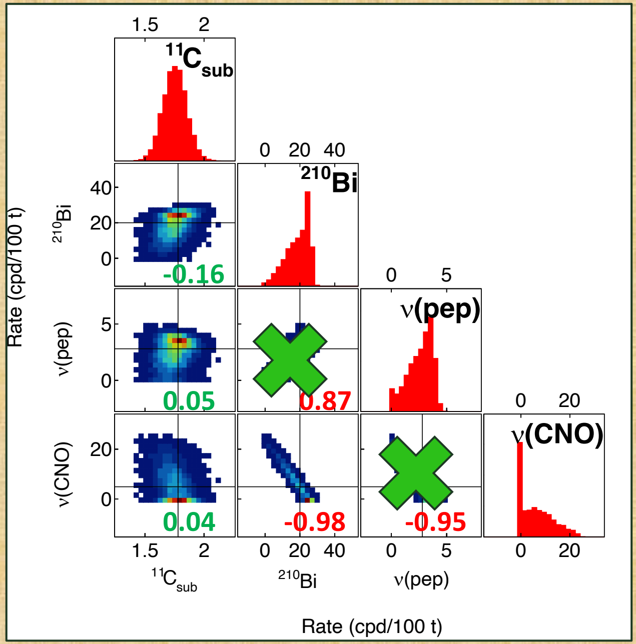


Expected rates:

- CNO ν ~ 4-5 cpd/100 ton
- pep ν ~ 3 cpd/100 ton
- ^{210}Bi ~ 15-20 cpd/100 ton

TOWARDS THE CNO-ν MEASUREMENT (2)

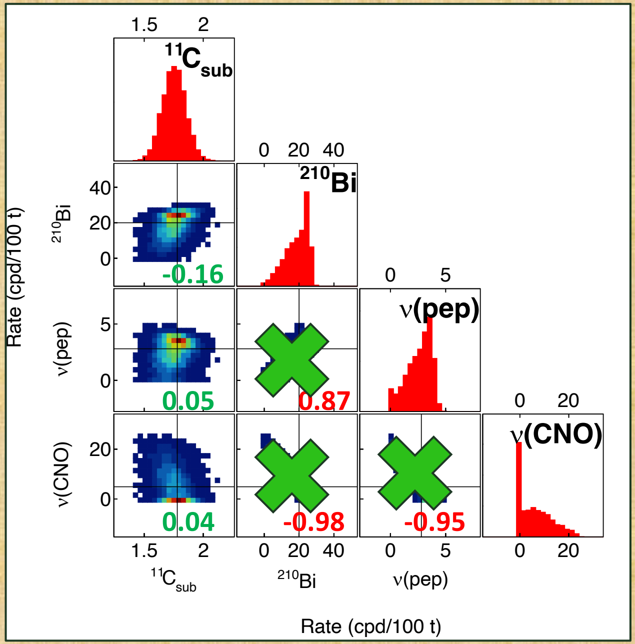
To reduce correlations we put a constraint on the pp/pep ratio following the theoretical predictions as described in *Nature 562 (2018) 505.*



The ^{210}Bi spectrum is still quasi-degenerate with the CNO neutrino one…..

TOWARDS THE CNO-ν MEASUREMENT (2)

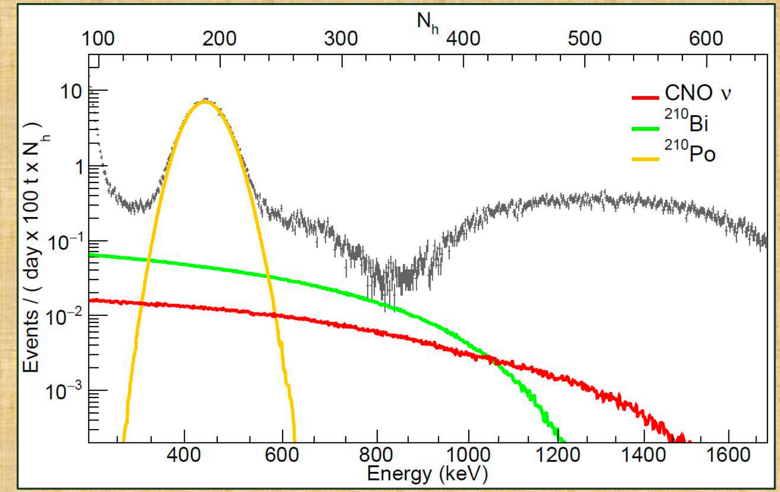
To reduce correlations we put a constraint on the pp/pep ratio following the theoretical predictions as described in *Nature 562 (2018) 505.*



The ^{210}Bi spectrum is still quasi-degenerate with the CNO neutrino one…..

…. But the ^{210}Bi rate can be constrained by precisely (and independently) mapping the ^{210}Po rate!

$$^{210}\mathrm{Pb} \xrightarrow[23\,y]{\beta^-} {}^{210}\mathrm{Bi} \xrightarrow[5\,d]{\beta^-} {}^{210}\mathrm{Po} \xrightarrow[138\,d]{\alpha} {}^{206}\mathrm{Pb}\ (\mathrm{stable})$$



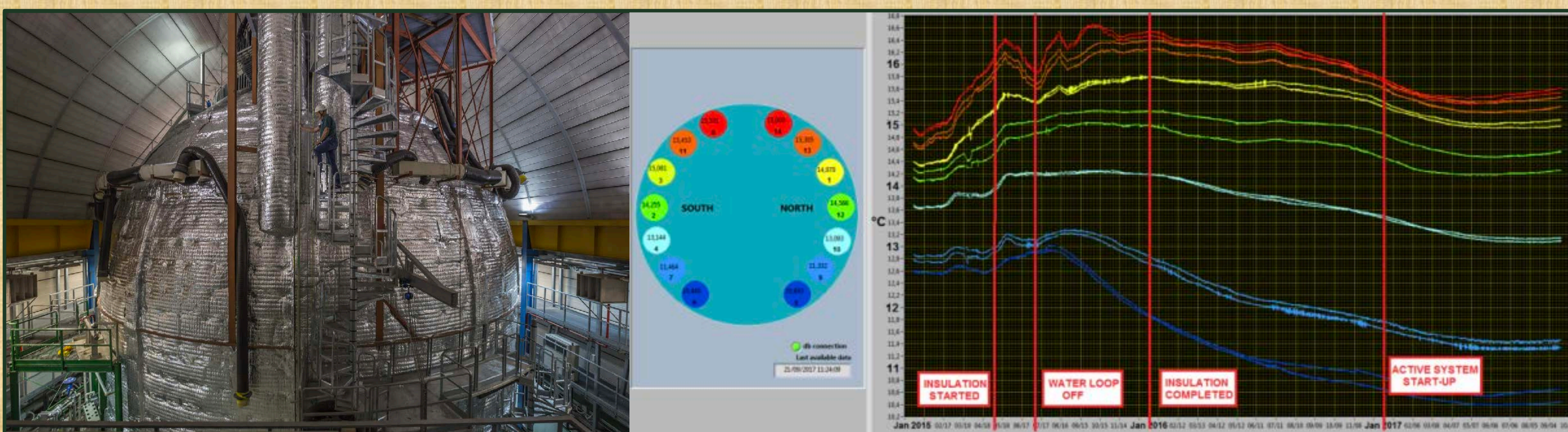
^{210}Po is "easier" to identify than ^{210}Bi:

- α decay → pulse shape discrimination
- Monoenergetic → "gaussian" peak

TOWARDS THE CNO MEASUREMENT (3)

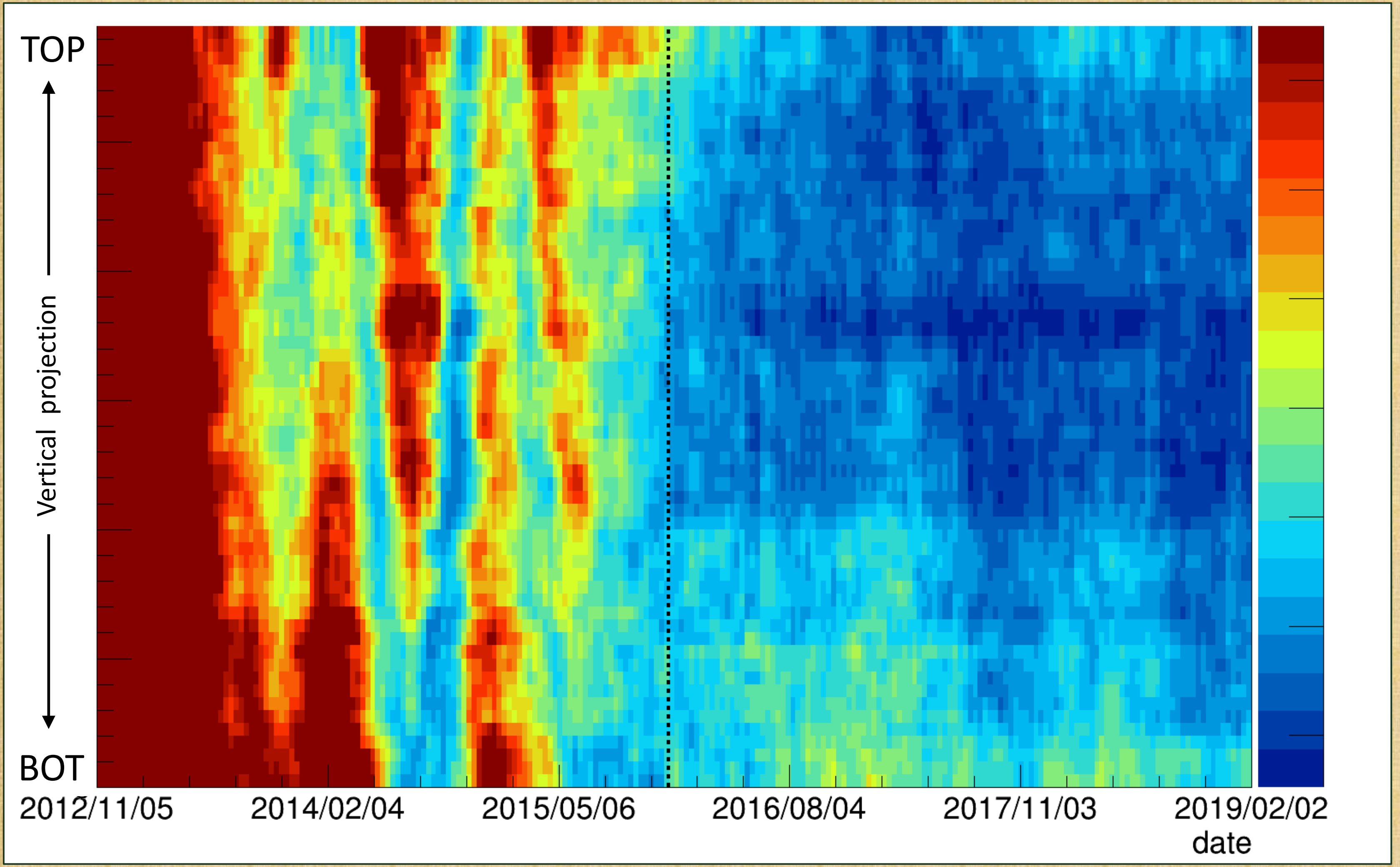
- Convective motions triggered by seasonal changes in temperature bring inside the scintillator an unknown amount of ^{210}Po which has been present on the nylon Inner Vessel.
- This breaks the secular equilibrium of the ^{210}Pb chain!

- **Before performing any counting analysis, we had to thermally insulate the detector to stop convective motions**! **DONE and COMPLETED in May 2016** ☺

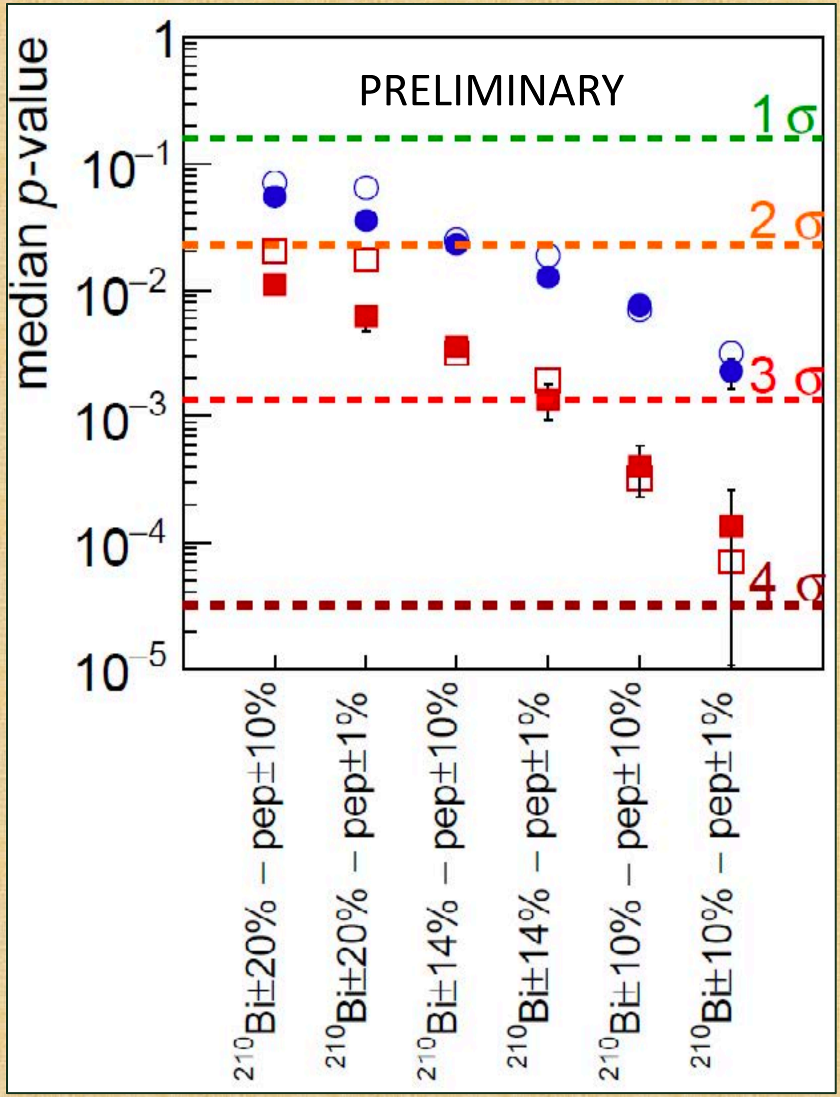


TOWARDS THE CNO MEASUREMENT (4)

^{210}Po counting rate inside the Inner Vessel scintillator volume



TOWARDS THE CNO MEASUREMENT (5)



Sensitivity studies

A sensitivity study has been performed to assess the detector performances in detecting CNO neutrinos assuming different constraints on ^{210}Bi and pep-ν rates, different metallicity scenario (HZ red, LZ blue) and different counting methods (filled/empty shapes)

CONCLUSIONS AND PERSPECTIVES

- Borexino has mapped out the entire pp solar fusion chain with high precision.

- Low-energy electron scattering can probe interesting new physics: we can simultaneously test the P_{ee} in the vacuum and matter dominated region.

- A new result on geoneutrinos has been released recently increasing considerably the measurement precision

- The collaboration is carrying on a huge effort with the goal of performing the first measurement of CNO neutrinos…stay tuned


Thank you for your attention!
The Borexino Collaboration